Prospect of Relativistic Laser Plasmas for Studying Fundamental Processes and Relativistic Astrophysics

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OUTLINE

1. Lasers and Astrophysics
2. Shock Waves
3. Reconnection of Magnetic Field Lines & Vortex Patterns
4. Relativistic Rotator
5. Flying Mirror for Femto-, Atto-, … Super Strong Fields
6. Overdense Accelerating Mirror
7. Conclusion

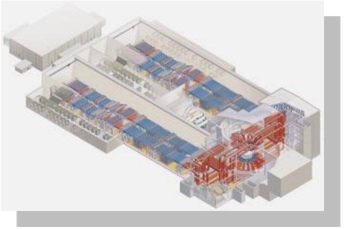
a) V. S. Berezinskii, S. V. Bulanov, V. L. Ginzburg, V. A. Dogiel, V. S. Ptuskin, Astrophysics of cosmic rays. (North Holland Publ.Co. Elsevier Sci. Publ. Amsterdam, 1990)

b) G.Mourou, T.Tajima, S.V.Bulanov, Optics in the Relativistic Regime, Rev. Mod. Phys. 78, 309 (2006)

1. Lasers and Astrophysics

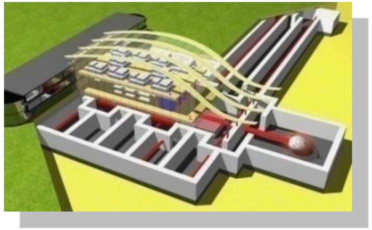
Morphology of Entities in Space and Laser Plasmas

NIF



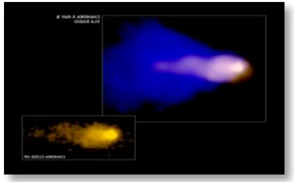
HiPER

ELI

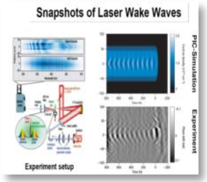


Wake

The Mouse Pulsar



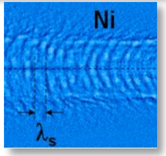
Electron Wake



Matlis et al, (2006)

Ion Wake

experiment



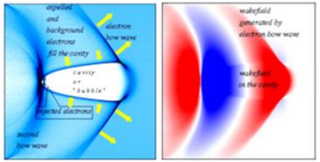
simulation

Borghesi, et al, (2005)

Bow Wave

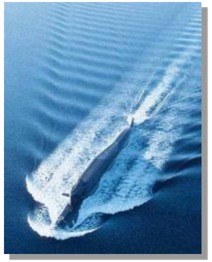
Chandra image of M87

Electron Bow Wave



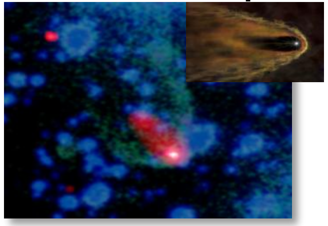
Esirkepov et al (2008)

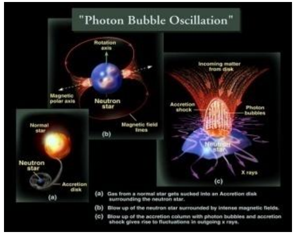
“Kalmar” Submarine



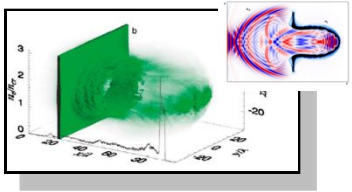
Photon Bubbles

“Black Widow” pulsar





RPDA



Esirkepov et al (2004)

Relativistic Limit in EM Wave – Plasma Interaction

Quiver energy of electron oscillating in the EM wave with the amplitude E_0 and frequency ω becomes larger than $m_e c^2$ when the dimensionless amplitude of the EM wave is greater than unity:

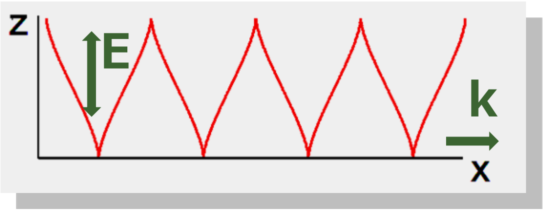
$$a_0 = \frac{eE_0}{m_e \omega c} > 1$$

In the EM wave interaction with the electron in vacuum its electron energy scales as (Landau & Lifshitz)

$$\mathcal{E} = \frac{1}{2} m_e c^2 a_0^2$$

When the electron oscillates in the EM wave propagating in a plasma we have (Akhiezer & Polovin)

$$\mathcal{E} = m_e c^2 a_0$$



Laser: Condition $a_0 > 1$ corresponds for 1μm laser wavelength to the intensity above 1.35×10^{18} W/cm^2

Today's lasers can provide the intensity $I > 2\times10^{22}$ W/cm^2 , i. e. $a_0 \approx 100$

V.Yanovskij et al (2008)

Magneto-dipole Radiation of Oblique Rotator

Space: Magneto-dipole radiation of oblique rotator, has been considered as a model for the pulsar radiation

Power emitted by rotator is given by $W = \frac{2}{3}\frac{\mu^2 (\sin\theta)^2 \omega^4}{c^3}$

Magnetic moment: $\mu \approx B r_p^3$; θ is the angle between $\vec{\mu}$ and $\vec{\omega}$

The EM wave intensity at the distance r is $I = W / 4\pi r^2$

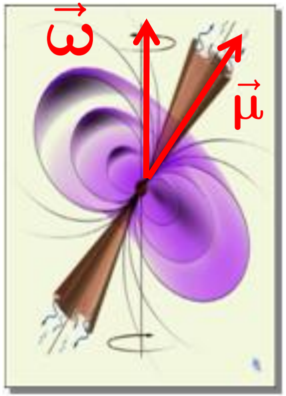
In the wave zone, $r = c/\omega$, the dimensionless wave amplitude is

$$a_0 = \frac{e\mu\omega^2}{m_e c^4}$$

For typical values of magnetic field, $B = 10^{12} G$, rotation frequency , $\omega = 200\, s^{-1}$,

and pulsar radius: $r_p = 10^6 cm$

it yields $\mu = 10^{30} G cm^3$ and . $a_0 \approx 10^{10}$



$r = \frac{c}{\omega}$

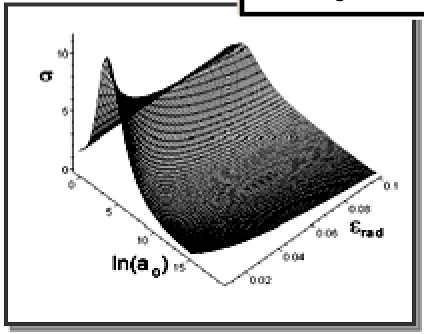
(Michel; Beskin, Gurevich, Istomin)

Crab pulsar

Amplitude $\left[a_0 = \frac{eE_0}{m_e c\omega}\right]$	Intensity $\left[\frac{\text{W}}{\text{cm}^2}\right]$	Regime
$a_{QED} = \frac{m_e c^2}{h\omega}$	2.4×10^{29}	e^+, e^- in vacuum
$a_{QM} = \frac{2e^2 m_e c}{3h^2\omega}$	5.6×10^{24}	quantum effects
$a_p = \frac{m_p}{m_e}$	1.3×10^{24}	relativistic p
$a_{rad} = \left(\frac{3\lambda}{4\pi r_e}\right)^{1/3}$	1×10^{23}	radiation damping
$a_{rel} = 1$	1.3×10^{18}	relativistic e^-

Ya. B. Zel'dovich, 1975; A.G.Zhidkov, et al., 2003
SVB, T. Zh. Esirkepov, J. Koga, T. Tajima, 2004

Cross section of nonlinear Thomson scattering



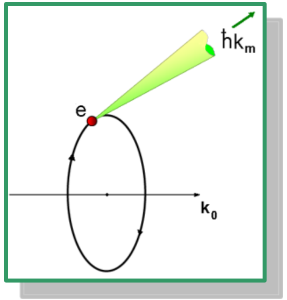
For the Crab pulsar,

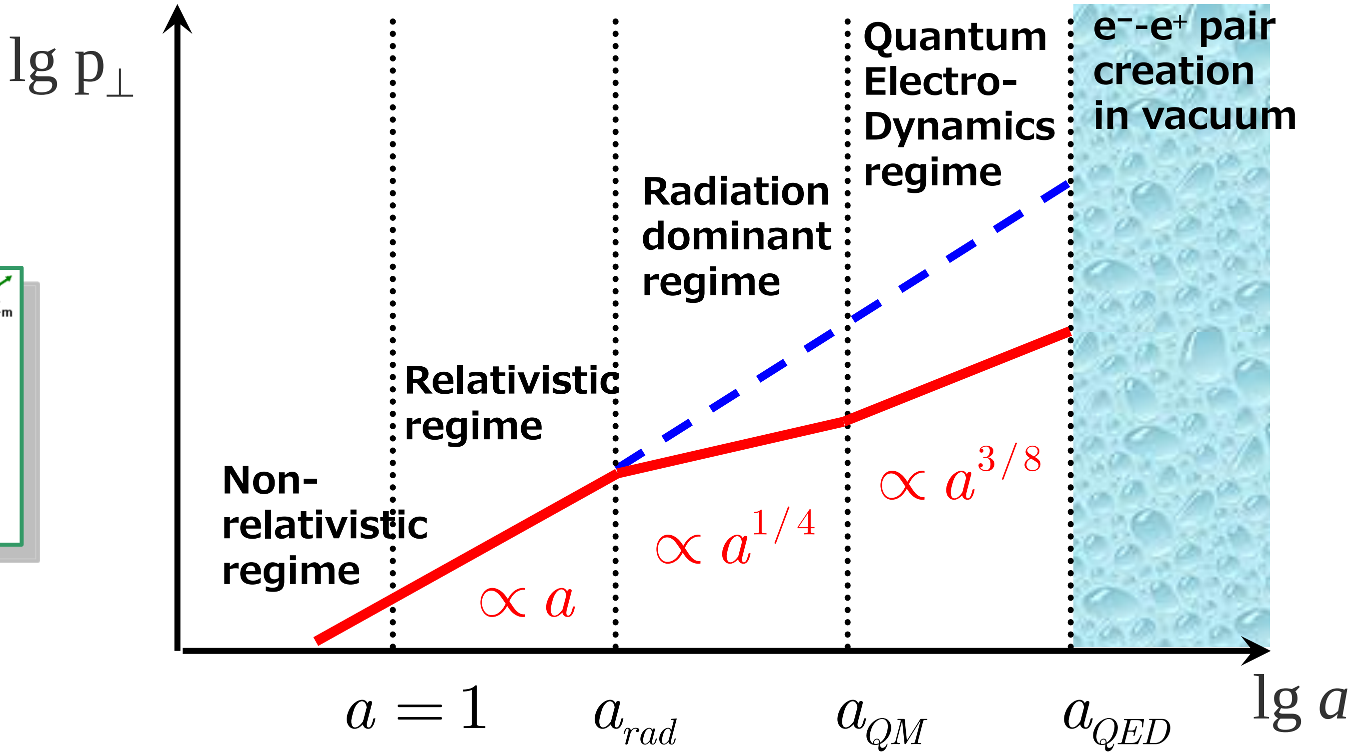
$$\omega = 200\, s^{-1}, \quad a_0 = 10^{10}$$

the radiation damping effects are crucially important because the EM wave amplitude is above the threshold:

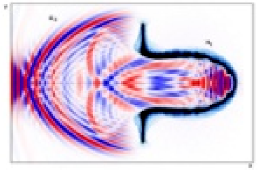
$$a_{rad} = \left(\frac{3\lambda}{4\pi r_e}\right)^{1/3} = 10^7$$

Laser-Plasma Interaction in the "Radiation-Dominant" Regime

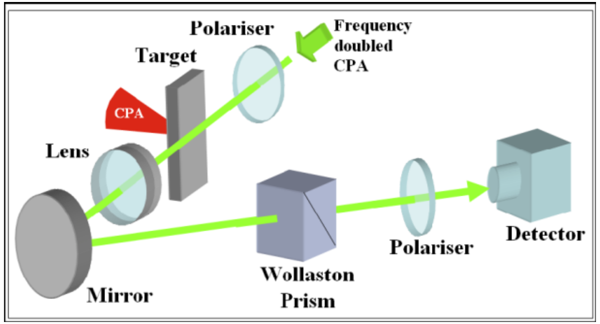


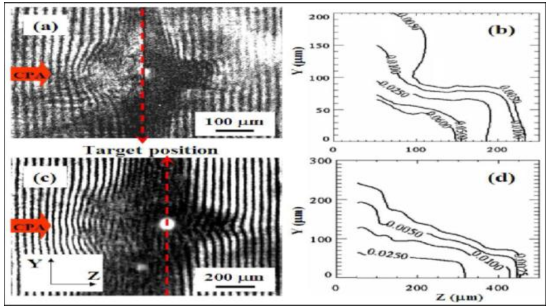


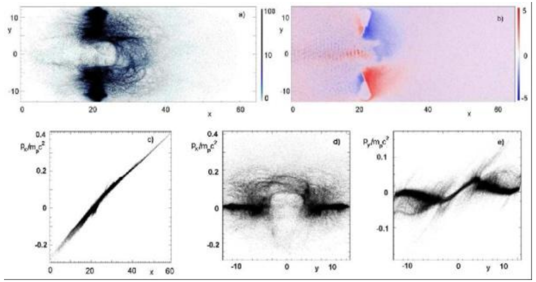
Plasma jets driven by Ultra-intense laser interaction with thin foils



VULCAN Nd-glass laser of Rutherford Appleton laboratory,
(60 J, 1ps & 250 J, 0.7 ps) interacts with foils (3, 5 mum, Al & Cu)





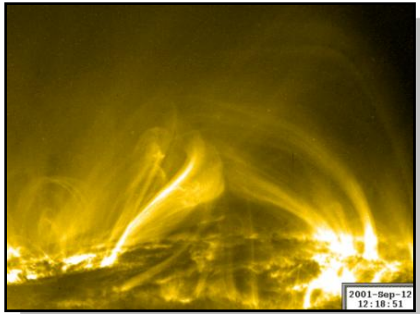


$$\frac{p^{(0)}}{m_p c} = \frac{2W(W+1)}{2W+1} \approx 2W$$

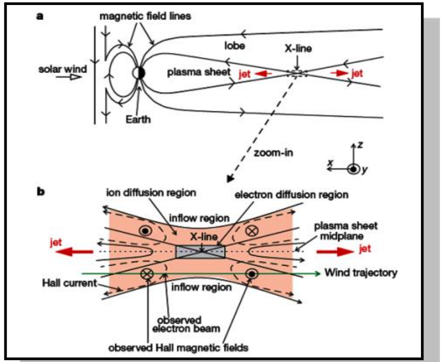
$$W = \int \frac{E^2(\psi)}{2\pi n_0 l} d\psi \ll 1$$

S. Kar, M. Borghesi, SVB, A.J. Mackinnon, P.K. Patel,
M.H. Key, L. Romagnani, A. Schiavi, A. Macchi, and O. Willi , PRL (2008)

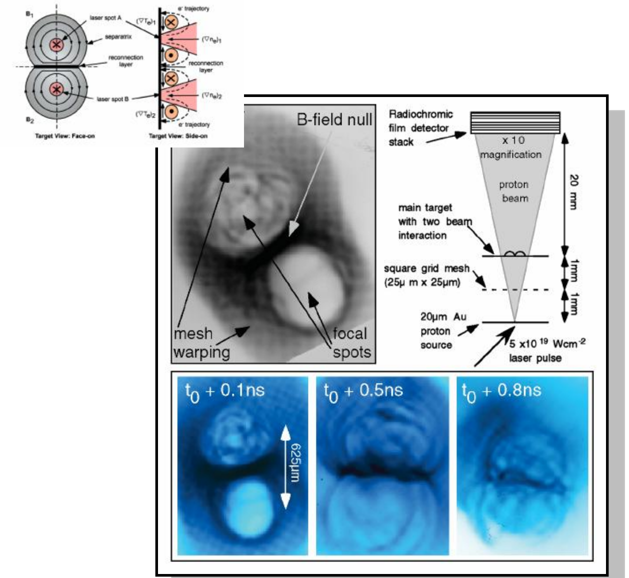
Reconnection of Magnetic Field Lines



Solar flare/ 12.09.2001(12:18:51)



M. Oieroset, et al., Nature (2001)



Nilson et al, PRL 97, 255001 (2006)

MAGNETIC RECONNECTION IN LASER PLASMAS HAS BEEN FORESEEN IN:

G.A.Askar'yan, SVB, F.Pegoraro, A.M.Pukhov,
Magnetic interaction of self-focused channels and magnetic wake excitation in high intense laser pulses,
Comments on Plasma Physics and Controlled Fusion 17, 35 (1995).

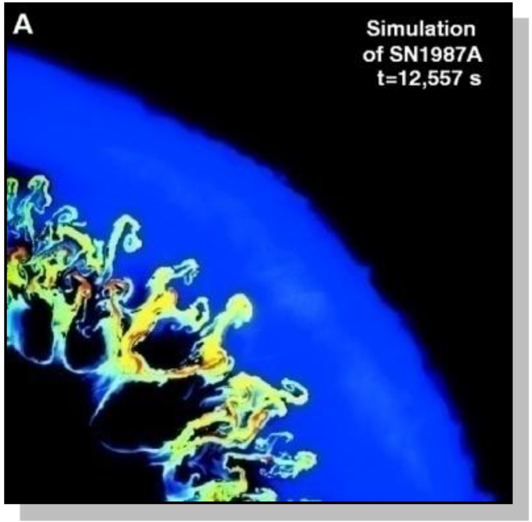
Laboratory Astrophysics

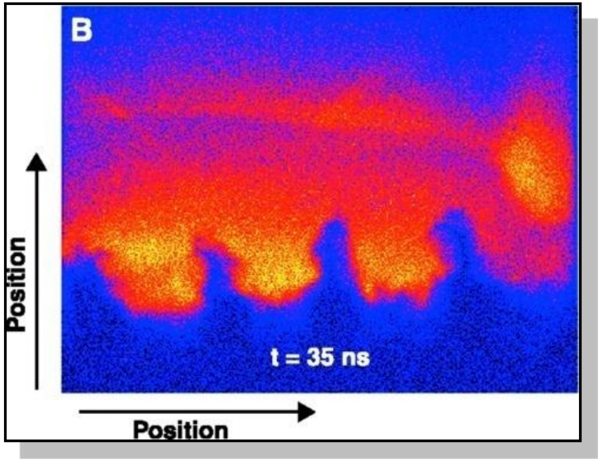
Laboratory Astrophysics

↓

Relativistic Laboratory Astrophysics with the Ultra Short Pulse High Power Lasers

We deal with the collisionless plasmas



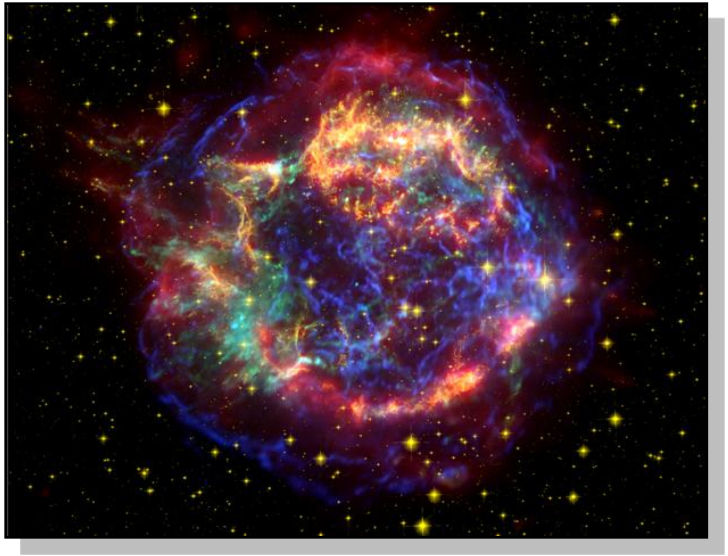


B. A. Remington et al., Science 284, 1488 (1999)

Rayleigh-Taylor & Richtmayer-Meshkov Instability,: seen in simulations of Supernovae (right) and in laser irradiated Nuclear Fusion target

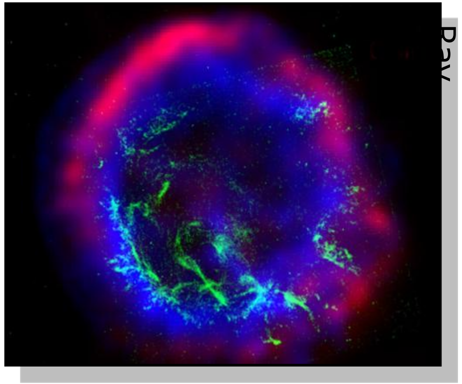
Radiative shock waves, plasma jets

2. Shock Waves



Cassiopea A

Shock Waves and RT Instability



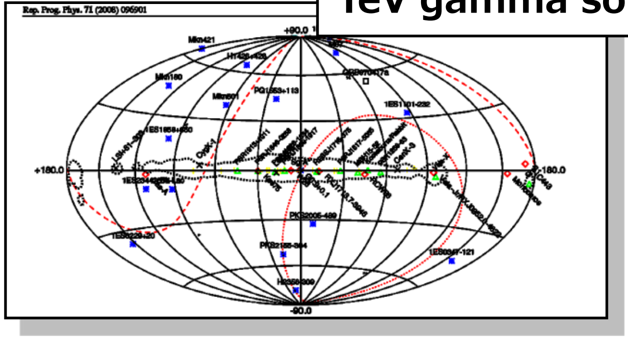
Supernova Remnant E0102-72 from Radio to X-Ray

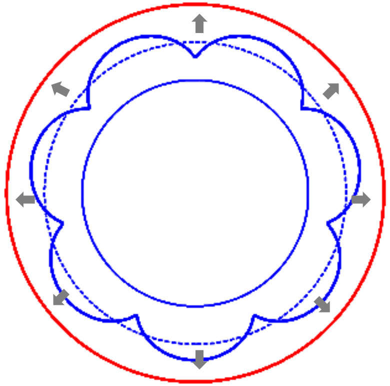
SNII $\mathcal{E}_{tot} = 10^{52} erg$

$1/10 \quad - \quad 1/30 \quad year$

$2\% \quad - \quad 10\%$

TeV gamma sources





1. Ballistic motion of the ejecta

2. Sedov's regime:

$$R_{SW} = 1.5(\mathcal{E}_{tot} t^2 / \rho)^{1/3} = \frac{5}{2} V_{SW} t$$

$$V_{SW} \propto t^{-3/5}$$

3. Radiation losses: $R_{SW} \propto t^{2/7}$

Collisionless Shock Waves

A structure of collisionless schock waves is determined by the counter play of dissipation and dispersion effects. These effects are described within the framework of the Korteweg-de Veries-Burgers equation:

$$\partial_t u + u\partial_x u - \nu\,\partial_{xx} u - \beta\,\partial_{xxx} u = 0$$

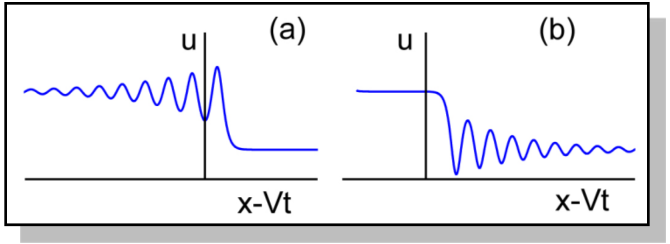
nonlinearity dispersion

dissipation

R.Z.Sagdeev, 1959

a) MS wave propagating almost perpendicularly to B field $\beta \approx v_a c^2 / 2\omega_{pe}$

with $v_a = B^2 / \sqrt{4\pi n m_p}$



b) MS wave propagation is almost parallel to B field

$\beta \approx -v_a c^2 / 2\omega_{pe}$

PRL **101**, 025004 (2008) PHYSICAL REVIEW LETTERS week ending 11 JULY 2008

Observation of Collisionless Shocks in Laser-Plasma Experiments

L. Romagnani,[1,*] S. V. Bulanov,[2,3] M. Borghesi,[1] P. Audebert,[4] J. C. Gauthier,[5] K. Löwenbrück,[6] A. J. Mackinnon,[7] P. Patel,[7] G. Pretzler,[6] T. Toncian,[6] and O. Willi[6]

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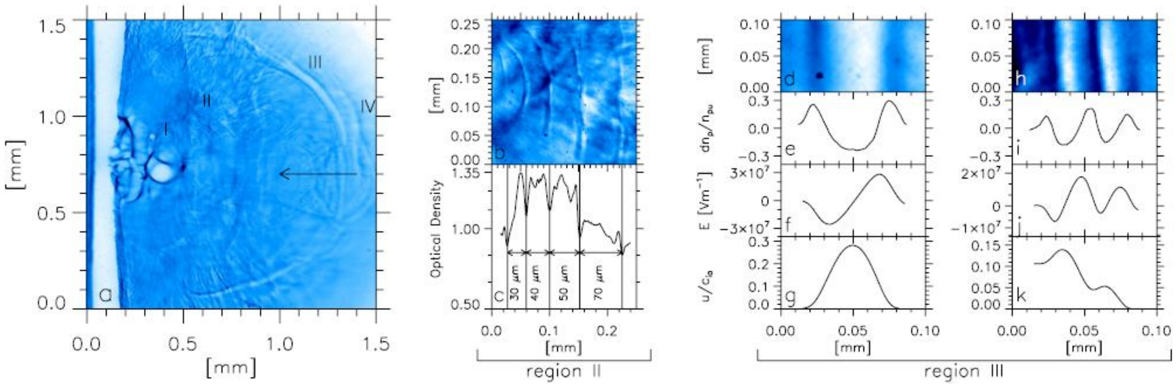
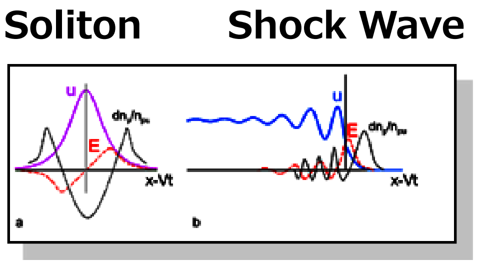
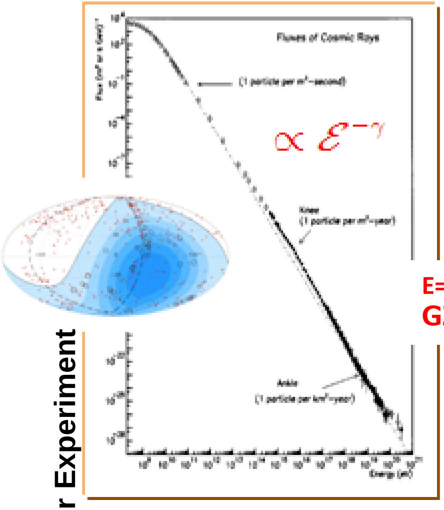
The propagation in a rarefied plasma ($n_e \lesssim 10^{15}$ cm^{-3}) of collisionless shock waves and ion-acoustic solitons, excited following the interaction of a long ($\tau_L \sim 470$ ps) and intense ($I \sim 10^{15}$ W cm^{-2}) laser pulse with solid targets, has been investigated via proton probing techniques. The shocks' structures and related electric field distributions were reconstructed with high spatial and temporal resolution. The experimental results were interpreted within the framework of the nonlinear wave description based on the Korteweg–de Vries–Burgers equation.





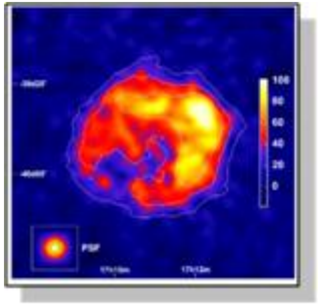
Acceleration at the Shock Wave Front

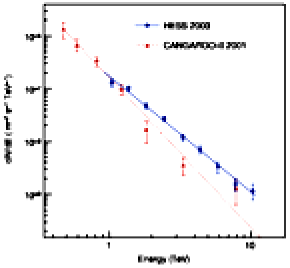
CR have a power law energy spectrum over several orders of magnitude energy range



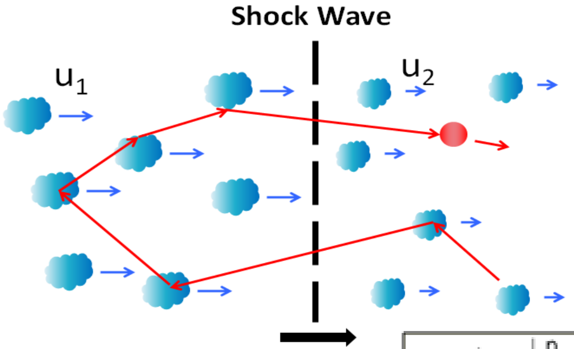
E=5x10^{19} eV
GZK-cutoff

The Auger Experiment



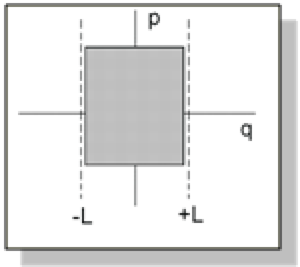


The gamma ray image of RX J1713.7-3946 spectrum and gamma ray obtained with the HESS telescope array (Aharonyan et al, 2008)



"Fermi acceleration"
G.F. Krymskii (1977)

$$\frac{\partial}{\partial x}\left(u(x)\,f - D\frac{\partial f}{\partial x}\right) = -\frac{2u_1}{3(\kappa+1)}\delta(x)\frac{1}{p^2}\frac{\partial}{\partial p}\left(p^2 f\right)$$



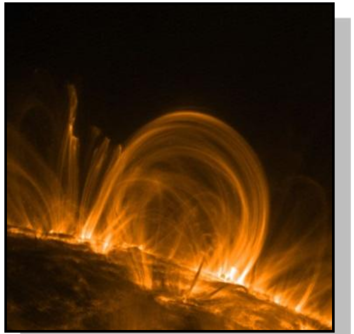
$$pL = const$$

$$u_2 = \frac{\kappa-1}{\kappa+1}u_1 \qquad u_1 > u_2$$

$$f(p) = Cp^{-\gamma}$$

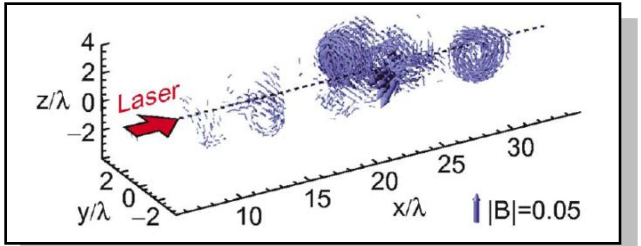
$$\gamma = \frac{3u_1}{u_1 - u_2}$$

3. Reconnection of Magnetic Field Lines & Vortex Patterns

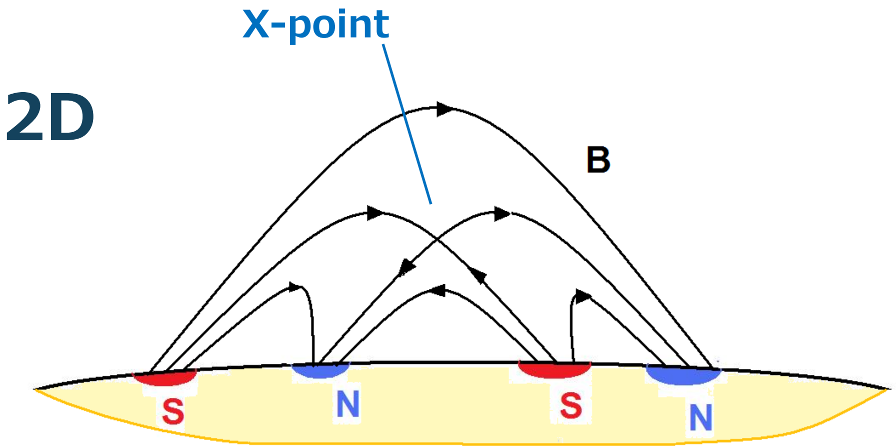


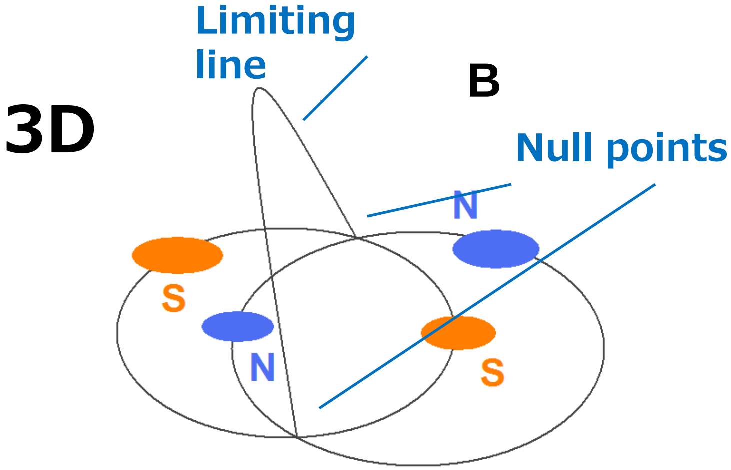
Solar Flare

Von Karman vortex row made by the wind over the Pacific island of Guadalupe

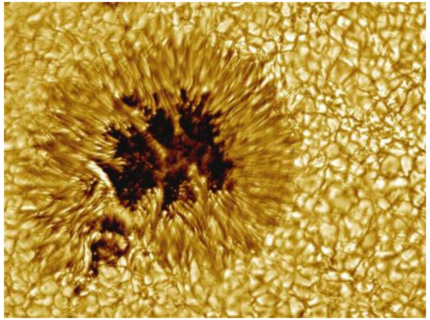


Magnetic (vortex) wake behind the laser pulse:
Esirkepov, et al., 2004

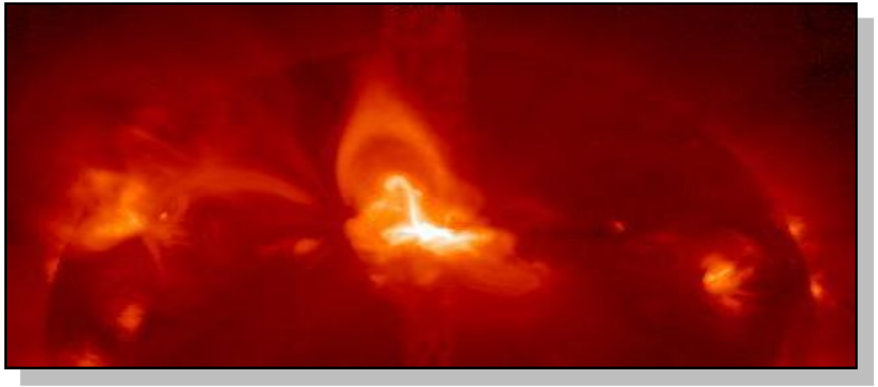




Sweet (1965)



Sunspot



Solar Flare

2D case: The field-line equation reads

$$\frac{dx}{B_x} = \frac{dy}{B_y} = ds$$

Using the relationships

$$B_x = \partial_y A_z - \partial_x F, \quad B_y = -\partial_x A_z - \partial_y F,$$

introducing complex variable $\varsigma = x + iy$, complex field and potential

$$B = B_x - iyB_y, \quad \Phi = F - iA_z,$$

we obtain the Hamiltonian equations for the magnetic field lines $(' = d/ds)$:

$$\varsigma' = -\frac{\partial \Phi}{\partial \varsigma}$$

The magnetic field lines are on the surfaces $A_z =$ constant

Local Structure of the Magnetic Field

Near null point we can expand the magnetic field as

$$\mathbf{B}(\mathbf{x},t) = (\mathbf{B}(0,t)\nabla)\mathbf{x} + \ldots.$$

Introducing the matrix $\partial B_i / \partial x_j \big|_{x=0} = A_{ij}, \quad B_i = A_{ij} x_j$

we write for the magnetic field lines $\dfrac{dx_i}{ds} = A_{ij} x_j.$

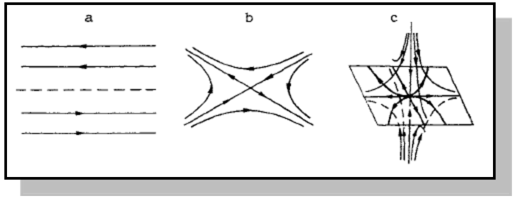
It yields $\det\left(A_{ij} - \lambda \delta_{ij}\right) = 0,$

The topology is determined by the eigenvalues $\lambda_\alpha \ \left(\sum_\alpha \lambda_\alpha = 0\right)$

We have the null surface, null line or null point depending on

$$\lambda_{1,2} = \pm\lambda' \text{ or } \lambda_{1,2} = \pm i\lambda'' \qquad \lambda_3 = 0$$

$$\lambda_{1,2} = \lambda' \pm i\lambda'' \qquad \lambda_3 = \lambda'$$



MHD Equations

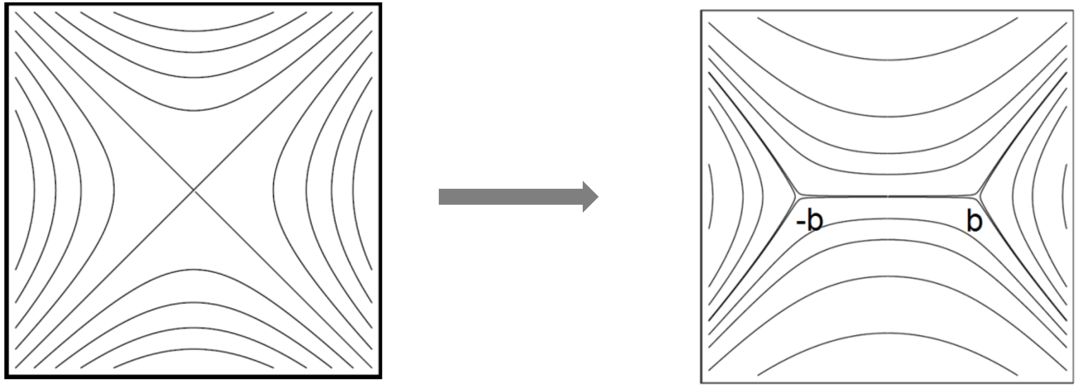
$$\partial_t \rho + \nabla(\rho \mathbf{v}) = 0,$$

$$\partial_t \mathbf{v} + (\mathbf{v}\nabla)\mathbf{v} = \frac{(\nabla \times \mathbf{B}) \times \mathbf{B}}{4\pi\rho},$$

$$\partial_t \mathbf{B} = \nabla \times (\mathbf{v} \times \mathbf{B}) + \nu_m \Delta \mathbf{B},$$

$$\nabla \cdot \mathbf{B} = 0$$

ρ - plasma density; $\mathbf{v}$ - velocity; $\mathbf{B}$ - magnetic field; ν_m - magnetic diffusivity



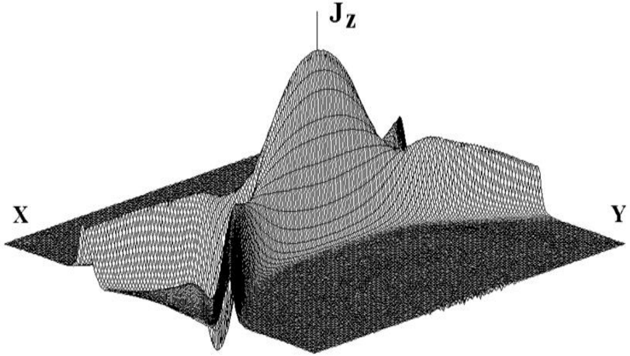
S. I. Syrovatskii, 1971

$$b = \sqrt{4I / hc}$$

$$\Phi = h\varsigma^2 / 2$$

$$\Phi = h\left[\varsigma\sqrt{\varsigma^2 - b^2} - \mathrm{Ln}\left(\varsigma - \sqrt{\varsigma^2 - b^2}\right)\right]$$

Current sheet near the X-line
of magnetic configuration



SVB, et al, 1996

Magnetic Reconnection in Collisionless Plasmas

In collisionless multispecies plasmas the **curl** of the canonical momentum

$$\mathbf{p}_\alpha = m_\alpha \mathbf{v}_\alpha + (e_\alpha / c)\mathbf{A}$$

is frozen in the corresponding flow velocity

$$\partial_t \nabla \times \mathbf{p}_\alpha = \nabla \times \left[\mathbf{v}_\alpha \times \nabla \times \mathbf{p}_\alpha\right]$$

The electron magnetohydrodynamics considers the dynamics of just the electrons, the ions are assumed to be at rest and the quasineutrality condition is fulfilled. The electron velocity is related to the magnetic field as

$$\mathbf{v}_e = -(c / 4\pi n_e)\nabla \times \mathbf{B}$$

with constant plasma density $n_e = n_i$. It yields

$$\partial_\mathbf{t}(\mathbf{B} - \Delta\mathbf{B}) = \nabla \times [(\nabla \times \mathbf{B}) \times (\mathbf{B} - \Delta\mathbf{B})]$$

In the linear approximation EMHD describes the whistler waves

The EMHD equations can be written as

$$\partial_t \mathbf{\Omega} = \nabla \times [(\nabla \times \mathbf{B}) \times \mathbf{\Omega}]$$

Here the generalized vorticity

$$\mathbf{\Omega} = \mathbf{B} - \Delta\mathbf{B} = \nabla \times (\mathbf{A} - \Delta\mathbf{A}) = \mathbf{B} + \nabla \times \mathbf{v}$$

is frozen into the electron fluid motion.

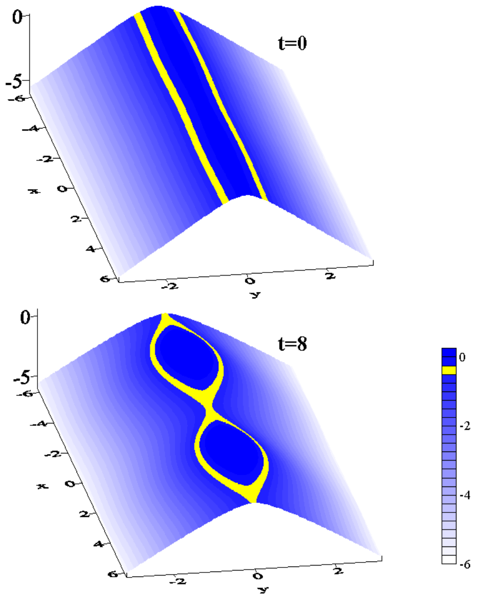
We consider the magnetic field given by

$$\mathbf{B} = \nabla \times (A_{||}\mathbf{e}_z) + B_{||}\mathbf{e}_z$$

The magnetic field pattern in the x, y plane is determined by

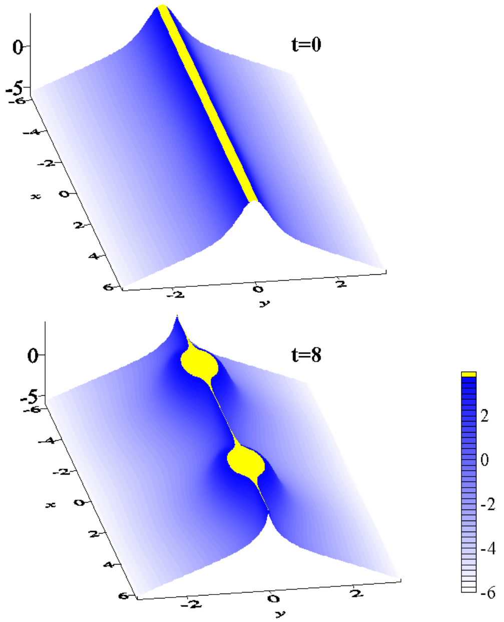
$$A_{||}(x, y, t) = const$$

Magnetic field



$$A_{||} = const$$

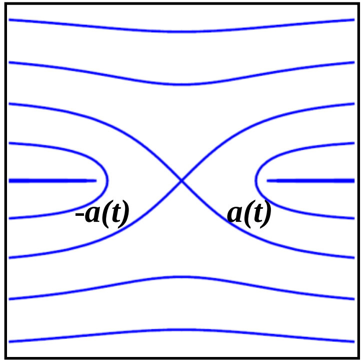
Generalized vorticity



$$A_{||} - \Delta A_{||} = const$$

K.Avinash, SVB, T.Esirkepov, P.Kaw, F.Pegoraro, P.Sasorov, A.Sen,
Forced Magnetic Field Line Reconnection in Electron Magnetohydrodynamics.
***Physics of Plasmas* 5, 2946 (1998)**

Charged Particle Acceleration



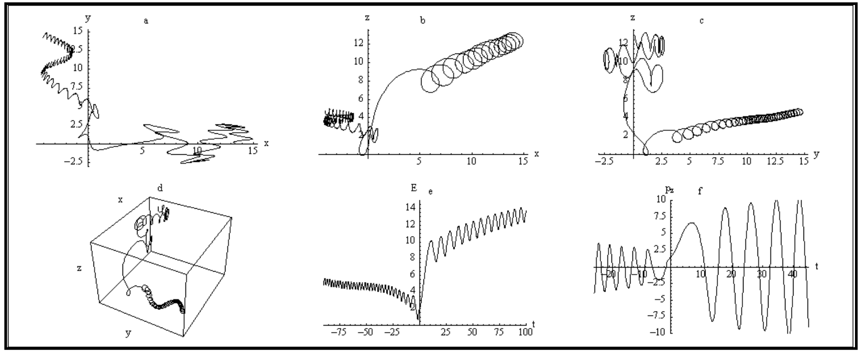
$$\Phi(\varsigma,t) = B_0\sqrt{a^2(t)-\varsigma^2}$$

In the vicinity of the X-line, the magnetic field is described by

$$B(\varsigma,t) = B_0\frac{\varsigma}{\sqrt{a^2(t)-\varsigma^2}} \approx B_0\frac{\varsigma}{a(t)}$$

and the electric field is given by

$$E(\varsigma,t) = -B_0\frac{a(t)\dot{a}(t)}{c\sqrt{a^2(t)-\varsigma^2}} \approx \frac{\dot{a}(t)}{c}B_0$$



The energy spectrum of fast particles is given by

$$\frac{d\mathcal{N}(\mathcal{E})}{d\mathcal{E}} \propto \exp\left(-\sqrt{\frac{2\mathcal{E}}{m\dot{a}^2}}\right)$$

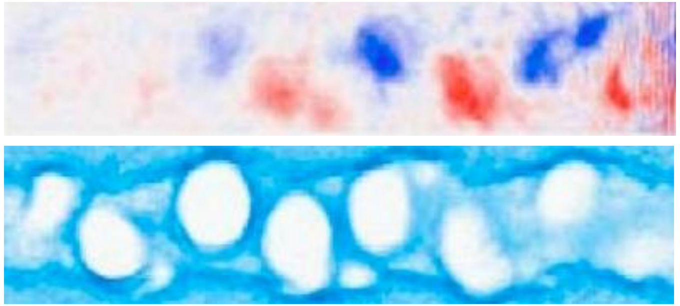
SVB, Sasorov; SVB, Cap

Electron Vortices behind the Laser Pulse

Antisymmetric vortex row

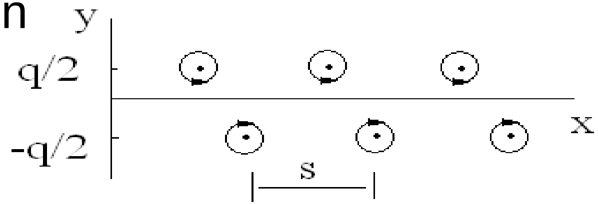
$B_z(x,y)$

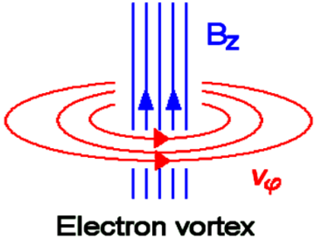
$n_i(x,y)$



L.M.Chen et al., Phys. Plasmas, 14, 040703 (2007)

Vortices described by the Hasegawa-Mima equation





Von Karman vortex row
H.Lamb, Hydrodynamics, 1947

SVB, T.Esirkepov, M.Lontano, F.Pegoraro,A.Pukhov, Phys. Rev. Letts. 76, 3562 (1996).

Interacting Point Vortices

As we know $\nabla\times(\mathbf{p}-e\mathbf{A}/c)$ is frozen:

$$(\partial_t+\mathbf{e}_z\times\nabla B\cdot\nabla)(\Delta B-B)=0$$

Discret vortices are described by equation $\Omega=\Delta B-B=\sum_j \Gamma_j\delta(\mathbf{r}-\mathbf{r}_j(t))$

its solution gives for the magnetic field $B=\sum_j B_j(\mathbf{r},\mathbf{r}_j(t))=-\sum_j\frac{\Gamma_j}{2\pi}K_0(|\mathbf{r}-\mathbf{r}_j(t)|)$

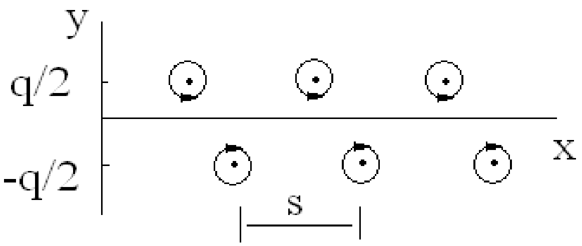
and the velocity of j-th vortex is

$$\frac{d\mathbf{r}_j}{dt}=\mathbf{e}_z\times\nabla\sum_{k\neq j}B_k(\mathbf{r}_j(t),\mathbf{r}_k(t))\Longleftrightarrow\begin{cases}\dfrac{dx_j}{dt}=\dfrac{1}{2\pi}\sum_{k\neq j}\Gamma_k\dfrac{y_k-y_j}{|\mathbf{r}_j(t)-\mathbf{r}_k(t)|}K_1(|\mathbf{r}_j(t)-\mathbf{r}_k(t)|)\\[2ex]\dfrac{dy_j}{dt}=\dfrac{1}{2\pi}\sum_{k\neq j}\Gamma_k\dfrac{x_j-x_k}{|\mathbf{r}_j(t)-\mathbf{r}_k(t)|}K_1(|\mathbf{r}_j(t)-\mathbf{r}_k(t)|)\end{cases}$$

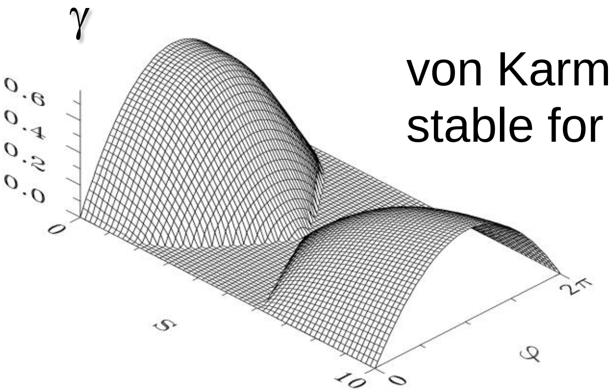
The Hamilton equations

Stability Domain

Antisymmetric vortex row

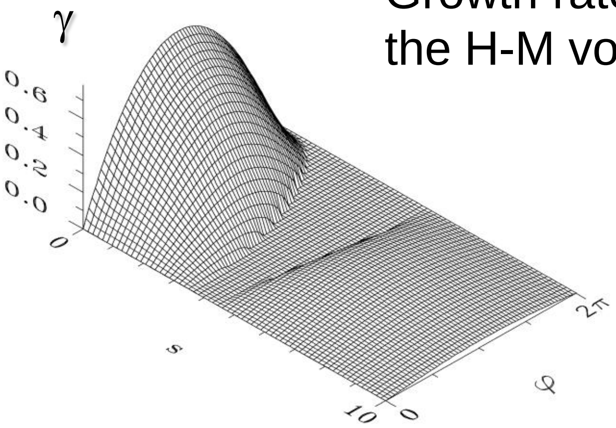


Lyapunov stability in the stability domain was proved

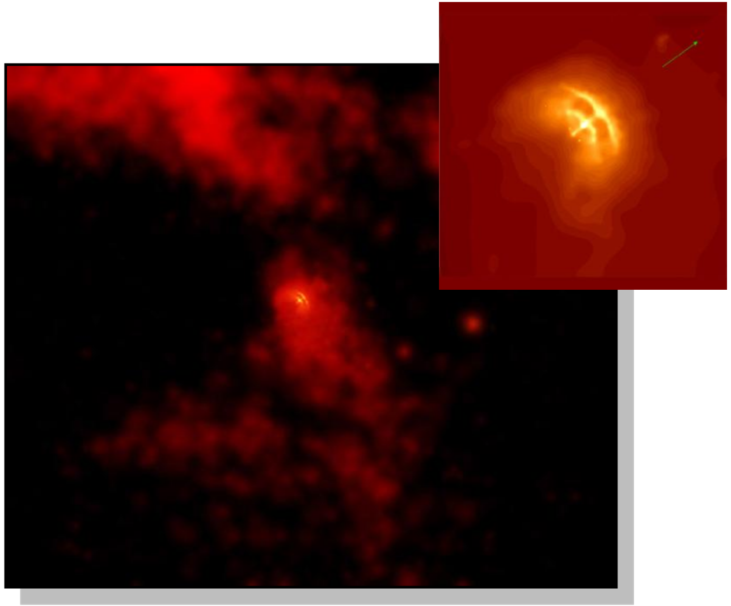


von Karman row is stable for q/s=0.281

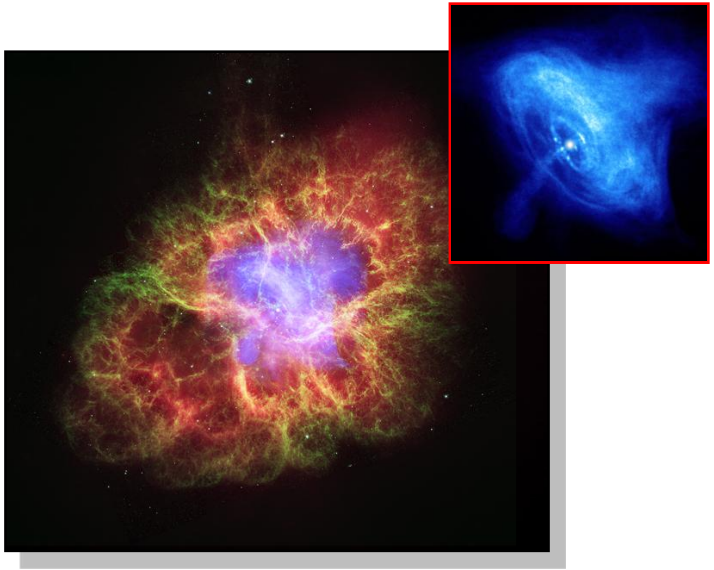
Growth rate vs q and s for the H-M vortex row

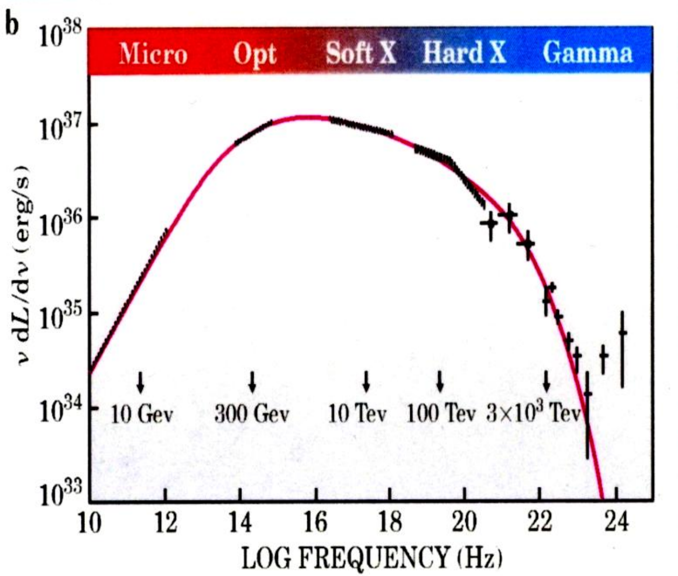


4. Relativistic Rotator



PeV γ from Crab Nebula





The Crab Pulsar, lies at the center of the Crab Nebula. The picture combines optical data (red) from the Hubble Space Telescope and x-ray images (blue) from the Chandra Observatory. The pulsar powers the x-ray and optical emission, accelerating charged particles and producing the x-rays.

ON THE PULSAR EMISSION MECHANISMS

1975

V. L. Ginzburg
P. N. Lebedev Physical Institute, Academy of Sciences of the USSR, Moscow, USSR

V. V. Zheleznyakov
Radio-Physical Institute, Gorkii, USSR

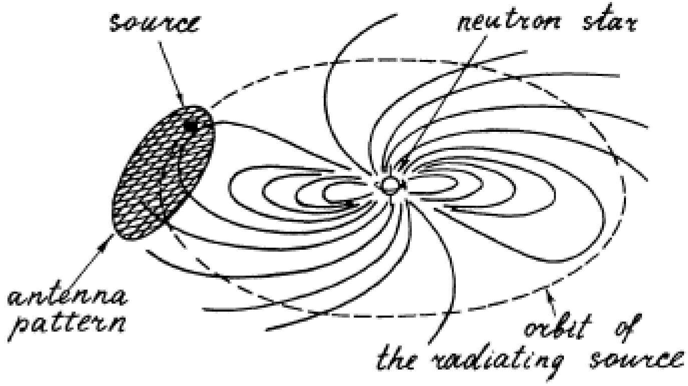
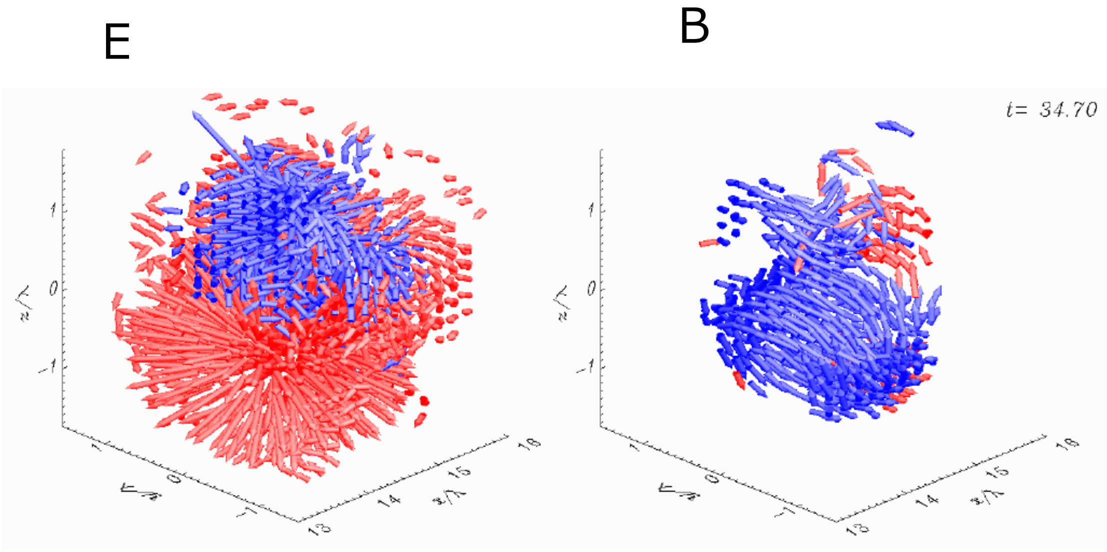
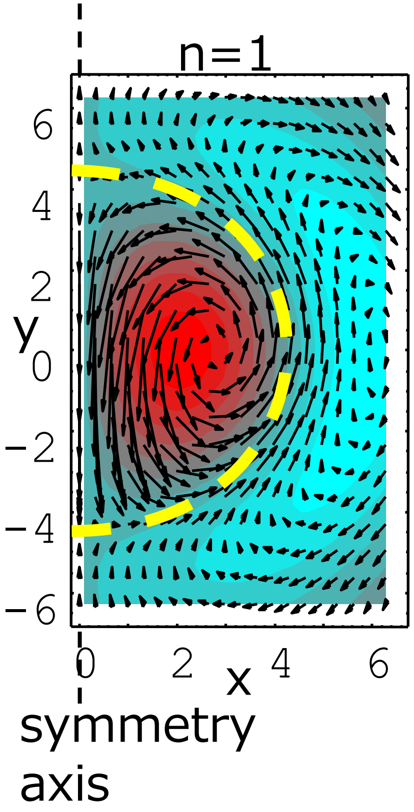


Figure 2 Schematic pulsar model.

Relativistic EM Soliton





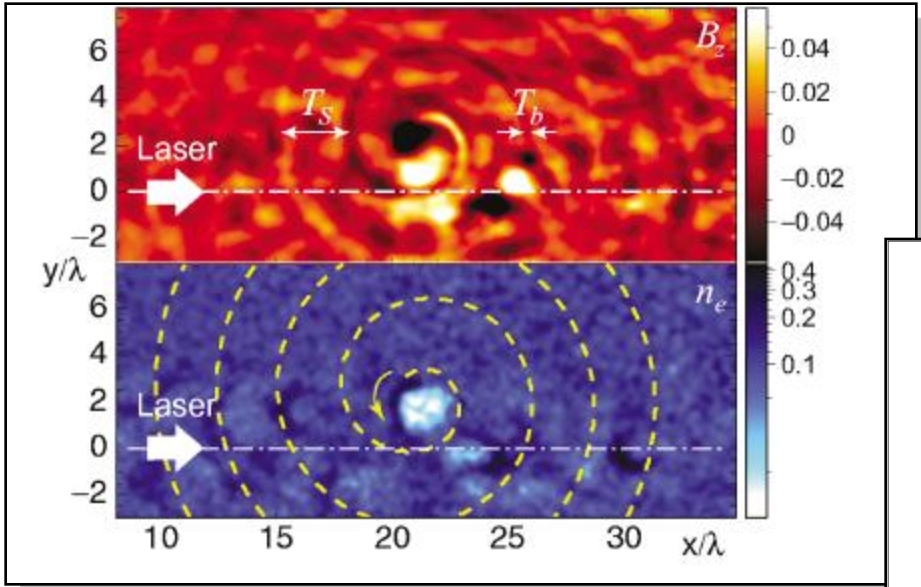
E.M. field in a spherical resonator

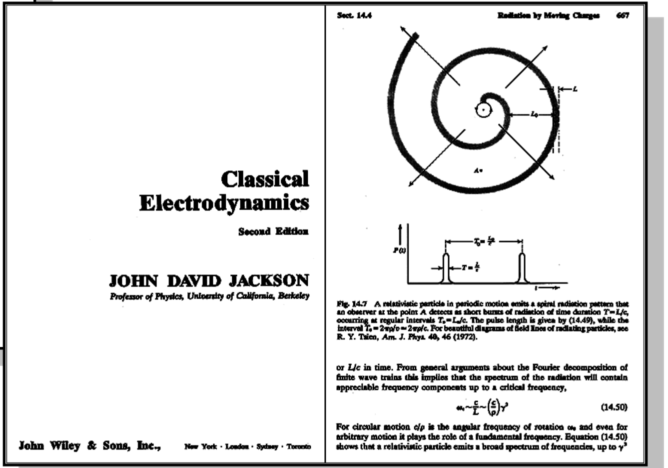
T.Esirkepov, SVB, K.Nishihara, F.Pegoraro, 2003

Circularly Polarized Soliton (3D PIC)

$B_z(x,y)$

$n_e(x,y)$





Classical Electrodynamics

Second Edition

JOHN DAVID JACKSON

Professor of Physics, University of California, Berkeley

John Wiley & Sons, Inc., New York · London · Sydney · Toronto

Sect. 14.4 Radiation by Moving Charges 667

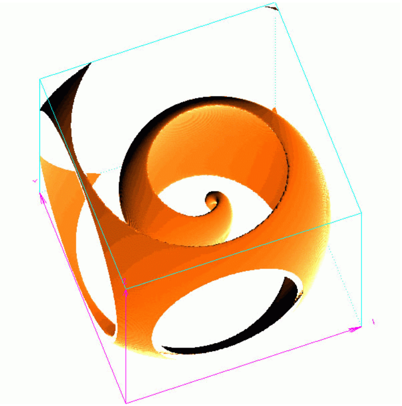
Fig. 14.7 A relativistic particle in periodic motion emits a spiral radiation pattern that an observer at the point *A* detects as short bursts of radiation of time duration $T = L/c$, occurring at regular intervals $T_b = L_0/c$. The pulse length is given by (14.49), while the interval $T_b = 2\pi\rho/v \approx 2\pi\rho/c$. For beautiful diagrams of field lines of radiating particles, see R. Y. Tsien, *Am. J. Phys.* 40, 46 (1972).

or L/c in time. From general arguments about the Fourier decomposition of finite wave trains this implies that the spectrum of the radiation will contain appreciable frequency components up to a critical frequency,

$$\omega_c \sim \frac{c}{L} \sim \left(\frac{c}{\rho}\right)\gamma^3 \qquad (14.50)$$

For circular motion c/ρ is the angular frequency of rotation ω_0 and even for arbitrary motion it plays the role of a fundamental frequency. Equation (14.50) shows that a relativistic particle emits a broad spectrum of frequencies, up to γ^3

T.Esirkepov, SVB, K.Nishihara, T.Tajima, 2004



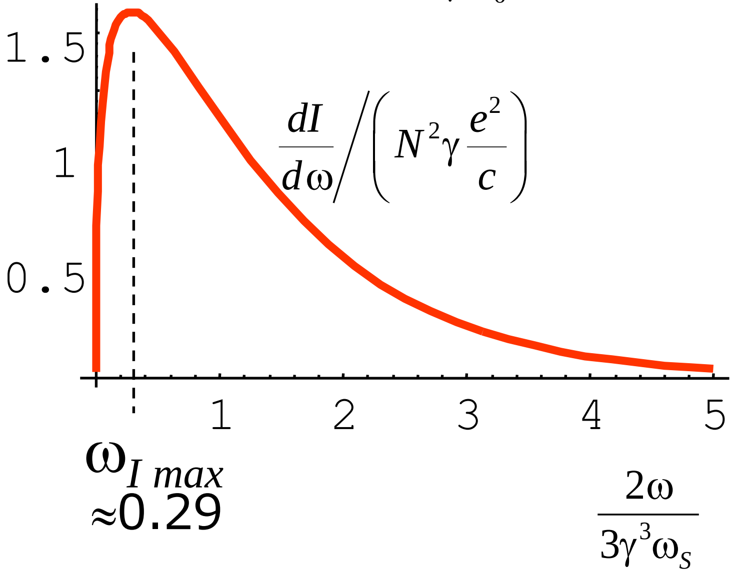
E.M. field energy density

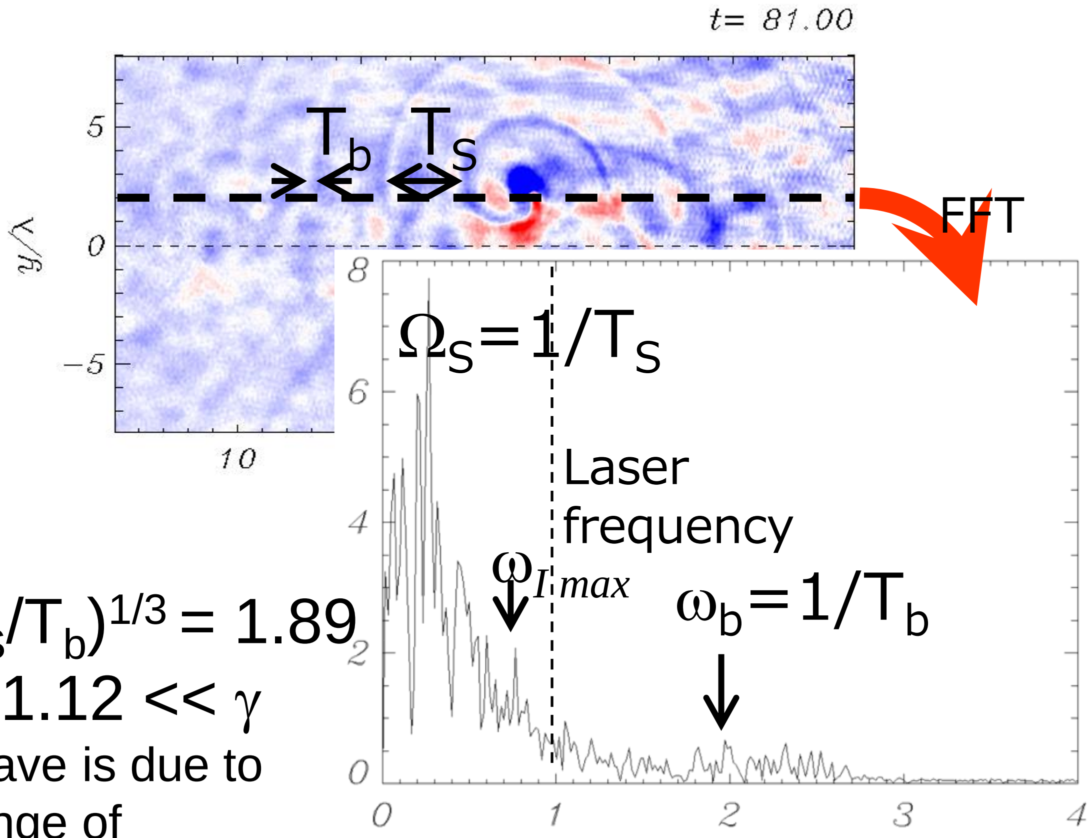
Energy loss by radiation

$$-\frac{d\mathcal{E}}{dt}=\frac{2e^2}{3c}N^2\omega_0^2\gamma^2\left(\gamma^2-1\right)$$

Frequency distribution of the total energy emitted by coherently rotating electrons

$$\frac{dI}{d\omega}=\sqrt{3}N^2\gamma\frac{e^2}{c}\frac{2\omega}{3\gamma^3\omega_0}\int_{\frac{2\omega}{3\gamma^3\omega_0}}^{\infty}\mathrm{K}_{5/3}(\xi)\,d\xi$$





$\gamma \approx (T_S/T_b)^{1/3} = 1.89$

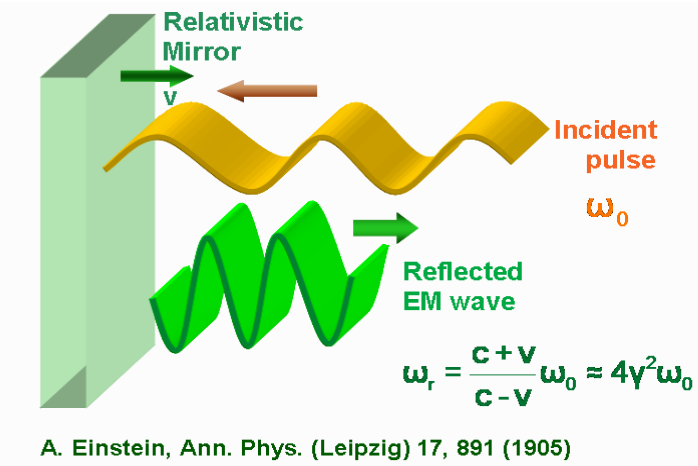
$\gamma_{e\ max}=1.12 << \gamma$

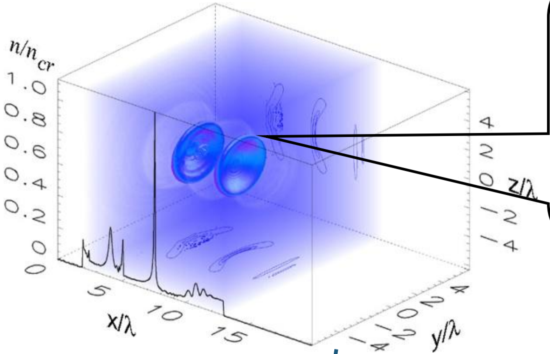
Spiral wave is due to fast change of electric charge density!

The solitons as the relativistic rotators can model the pulsar radiation under the earth laboratory conditions

5. Flying Mirror
for Femto-, Atto-, …
Super Strong Field Science

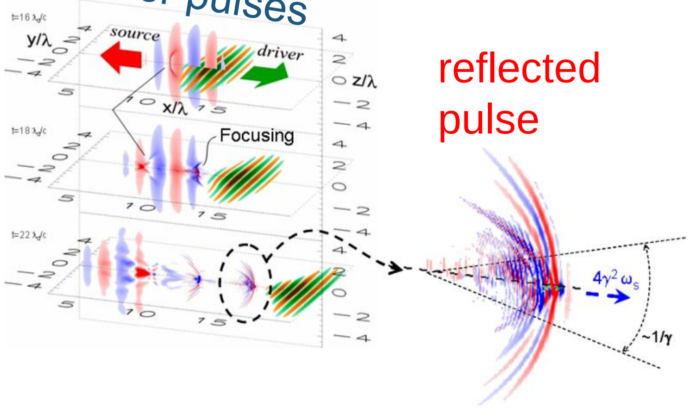
Flying Mirror Concept





paraboloidal relativistic mirrors formed by the wake wave left behind the laser driver pulse

laser pulses



reflected pulse

Frequency up-shifting and intensification of the light reflected at the relativistic mirror

S. Bulanov, T. Esirkepov, T. Tajima, Phys. Rev. Lett. 91, 085001 (2003)

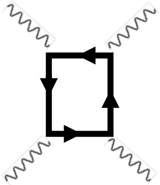
Laser Energy & Power to Achieve the Schwinger Field

The driver and source must carry 10 kJ and 30 J, respectively

Reflected intensity can approach the Schwinger limit $I_{QED} = 10^{29} W / cm^2$

$$E_{QED} = \frac{m_e^2 c^3}{e\hbar}$$

It becomes possible to investigate such the fundamental problems of nowadays physics, as e.g. the electron-positron pair creation in vacuum and the photon-photon scattering



$$\mathcal{L} = \frac{1}{16\pi} F_{\alpha\beta} F^{\alpha\beta} - \frac{\kappa}{64\pi}\left[5\left(F_{\alpha\beta}F^{\alpha\beta}\right)^2 - 14 F_{\alpha\beta}F^{\beta\gamma}F_{\gamma\delta}F^{\delta\mu}\right]$$

The critical power for nonlinear vacuum effects is $\mathcal{P}_{cr} = \frac{45\pi^2}{\alpha}\frac{cE_{QED}^2\lambda^2}{4\pi}$

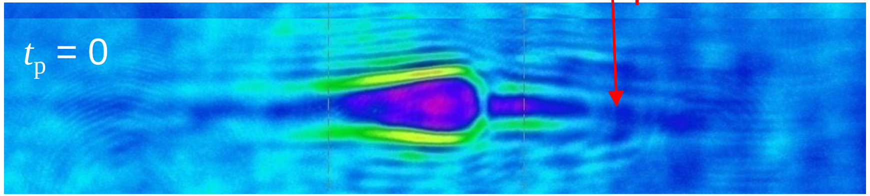
$$\mathcal{P}_{cr} \approx 2.5 \times 10^{24} W$$

Light compression and focusing with the FLYING MIRRORS yields $\mathcal{P} = \mathcal{P}_0 \gamma_{ph}$

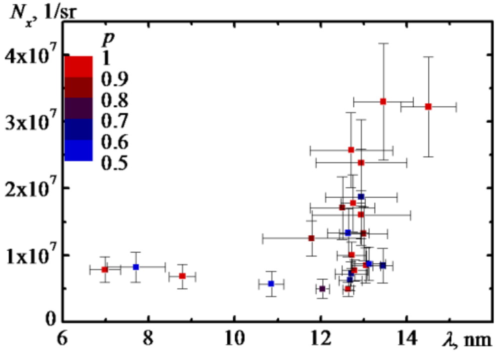
for $\lambda_0 = 1\mu m$ $\lambda = \lambda_0 / 4\gamma_{ph}^2$ **with** $\gamma_{ph} \approx 30$ **the driver power** **P_{cr}=10 PW**

Proof of Principle Experiment

In our experiments, narrow band XUV generation was demonstrated

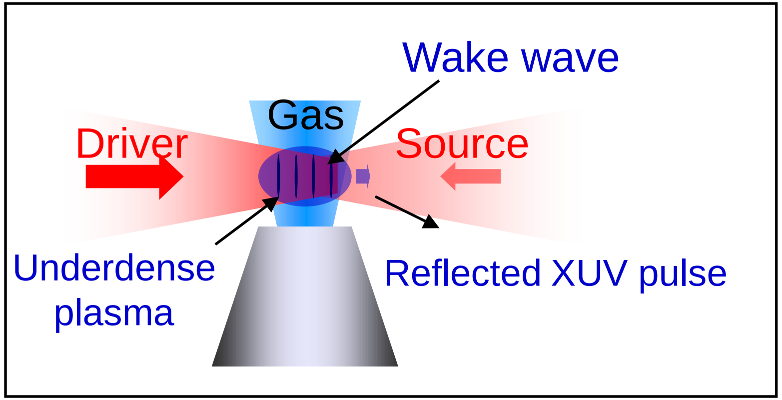


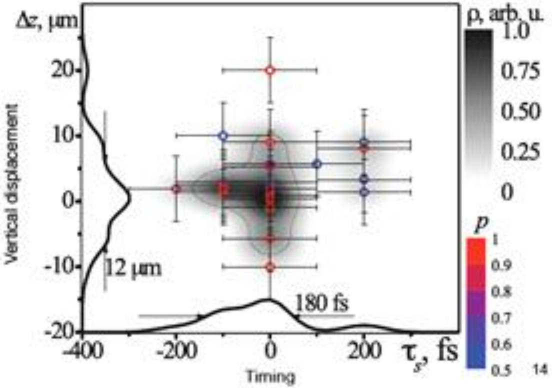
Detected signals



Frequency multiplication

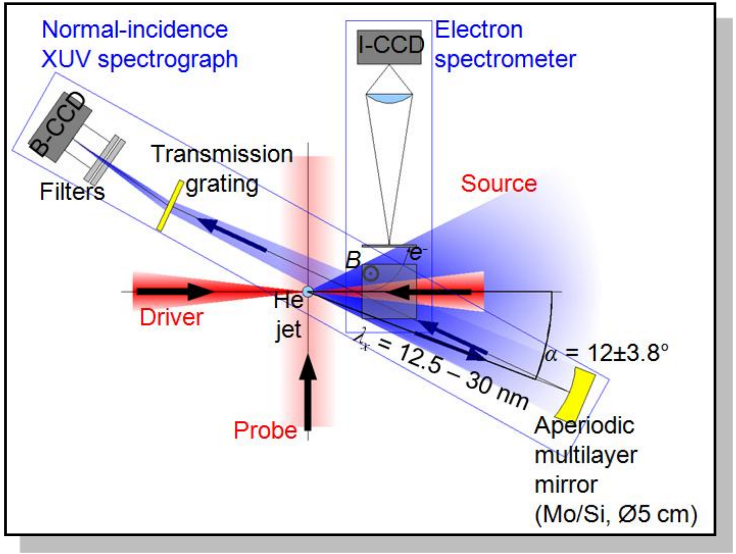
$$\frac{\omega_r}{\omega_s} = 55\ldots114$$

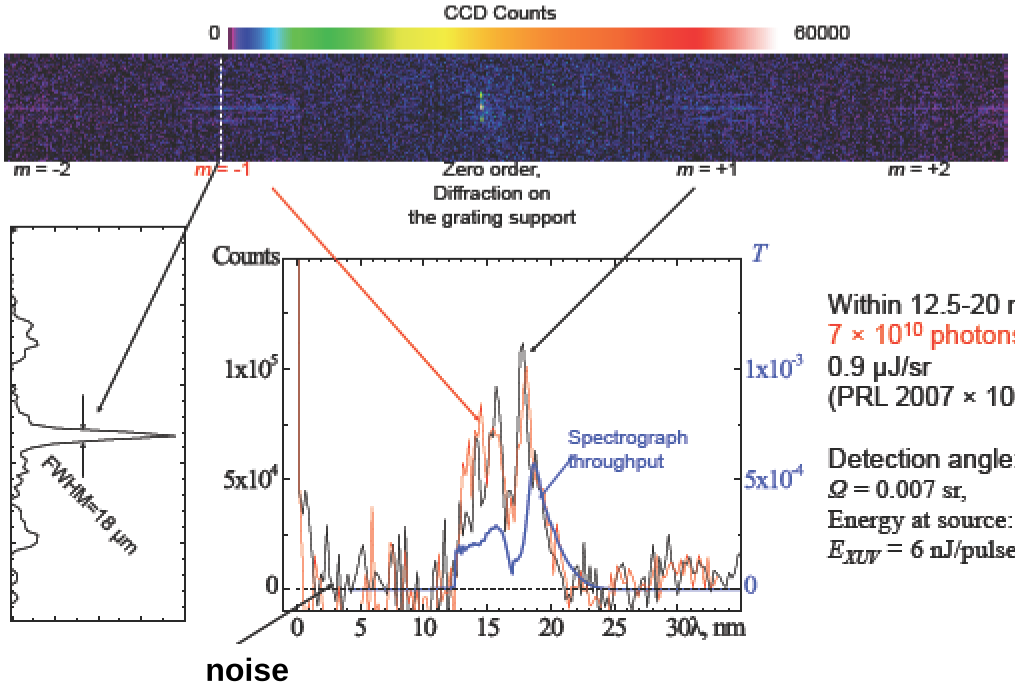




M. Kando, et al., Phys. Rev. Lett. 99, 135001 (2007);
A. Pirozhkov, et al., Phys. Plasmas 14, 123106 (2007)

Flying Mirror in the Head-On Collision Experiment



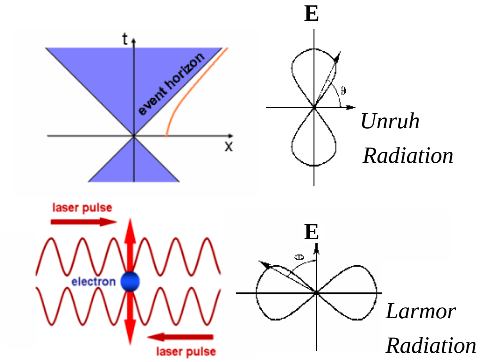


Within 12.5-20 nm:
7 × 10^{10} photons/sr
0.9 µJ/sr
(PRL 2007 × 1000)

Detection angle:
Ω = 0.007 sr,
Energy at source:
E_{XUV} = 6 nJ/pulse

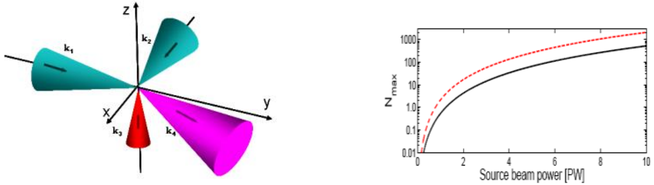
Two head-on colliding laser pulses

High Field Science

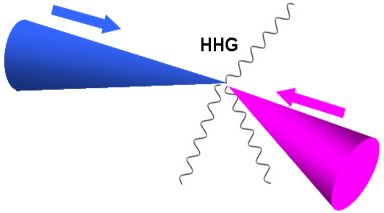


Unruh radiation (Chen&Tajima (1999))

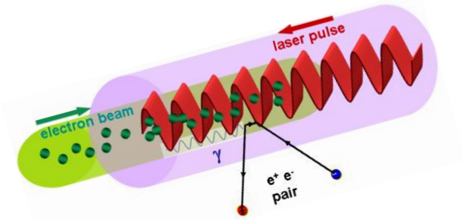
Birefringent e.m. vacuum (Rozanov (1993))



4-wave mixing (Lundström et al (2006))



Higher harmonic generation through quantum vacuum interaction (Fedotov & Narozhny (2006); Di Piazza, Keitel)



Electron-positron pair production in the laser interaction with the electron beam: $e^- + n\gamma \rightarrow \gamma,\ \gamma + n\gamma' \rightarrow e^+ + e^-$
Bula et al (1996); Burke et al (1997)

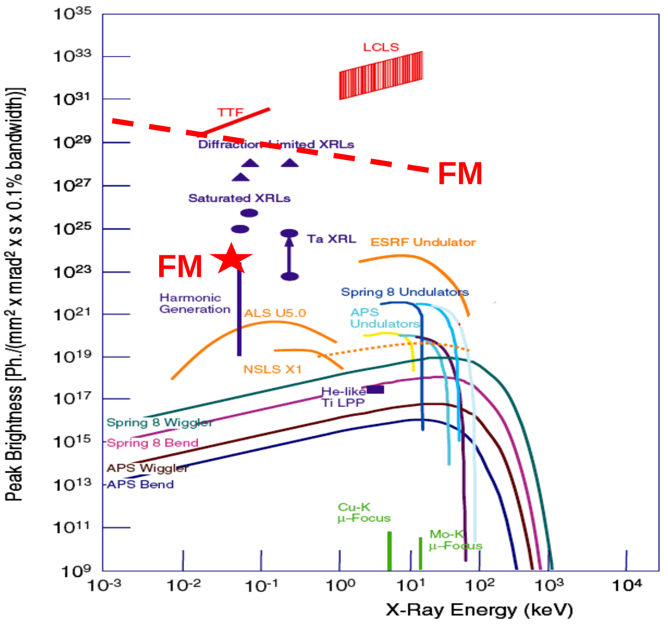
Compact Coherent Ultrafast X-Ray Source

X-ray source	Wavelength	Pulse Duration	Pulse Energy	Mono-chromaticity (Δλ/λ)	Coherence
XFEL (DESY)	13.8 nm	50 fs	100 μJ	10^{-3}	spatial good
Plasma XRL	13.9 nm	7 ps	10 μJ	10^{-4}	spatial good
Laser plasma	wide spectrum 1 nm – 40 nm	1 ps – 1 ns	10 μJ	10^{-2} – 10^{-3}	No
HHG	5 – 200 nm	100 attosec	1 μJ	10^{-2} – 10^{-3}	spatial and temporal good
Flying Mirror	0.1 – 20 nm	< 1 fs	1 mJ	10^{-2} – 10^{-4}	spatial and temporal good

Predicted by the FM theory parameters of the x-ray pulse compared with the parameters of high power x-ray generated by other sources

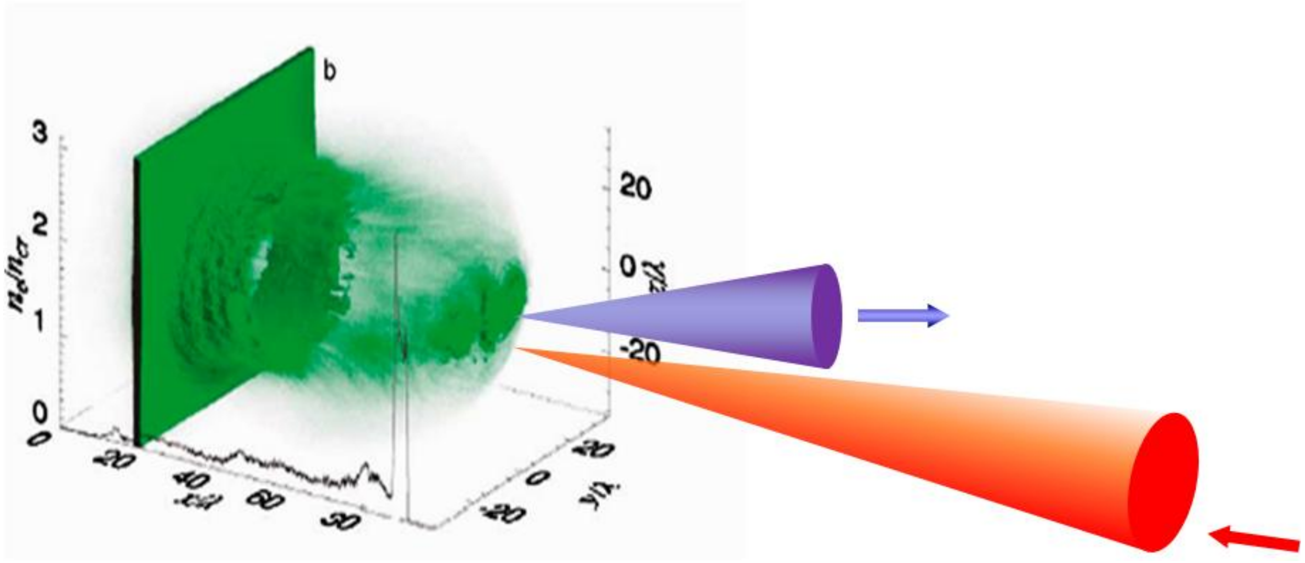
Brightness

$$B = 2\times10^{28}\left(\frac{\mathcal{E}_{las}}{1\,\mathrm{J}}\right)\sqrt{\frac{1\,\mathrm{KeV}}{\hbar\omega_\gamma}}\,\frac{1}{\mathrm{mm^2\,mrad^2\,0.1\%\,bandwidth}}$$

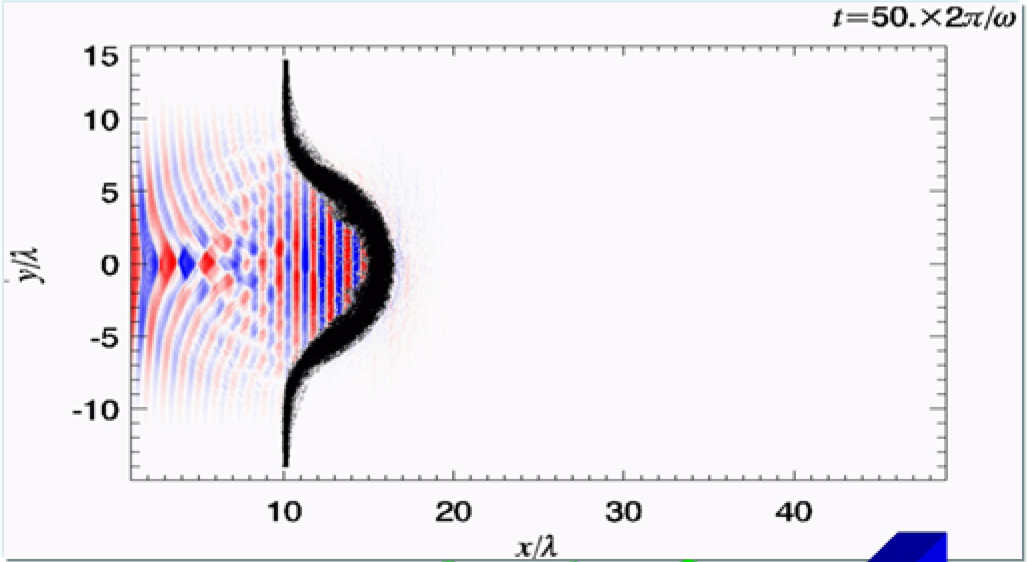


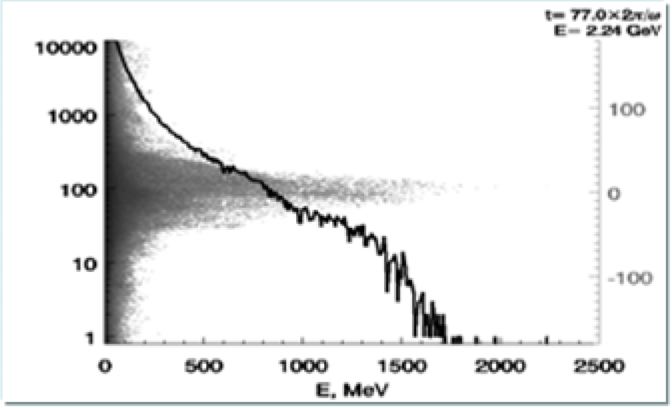
Peak brightness of various light sources

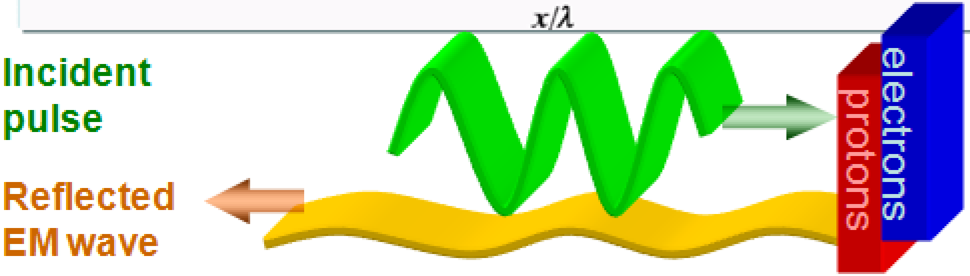
6. Overdense Accelerating Mirror



Accelerating mirror

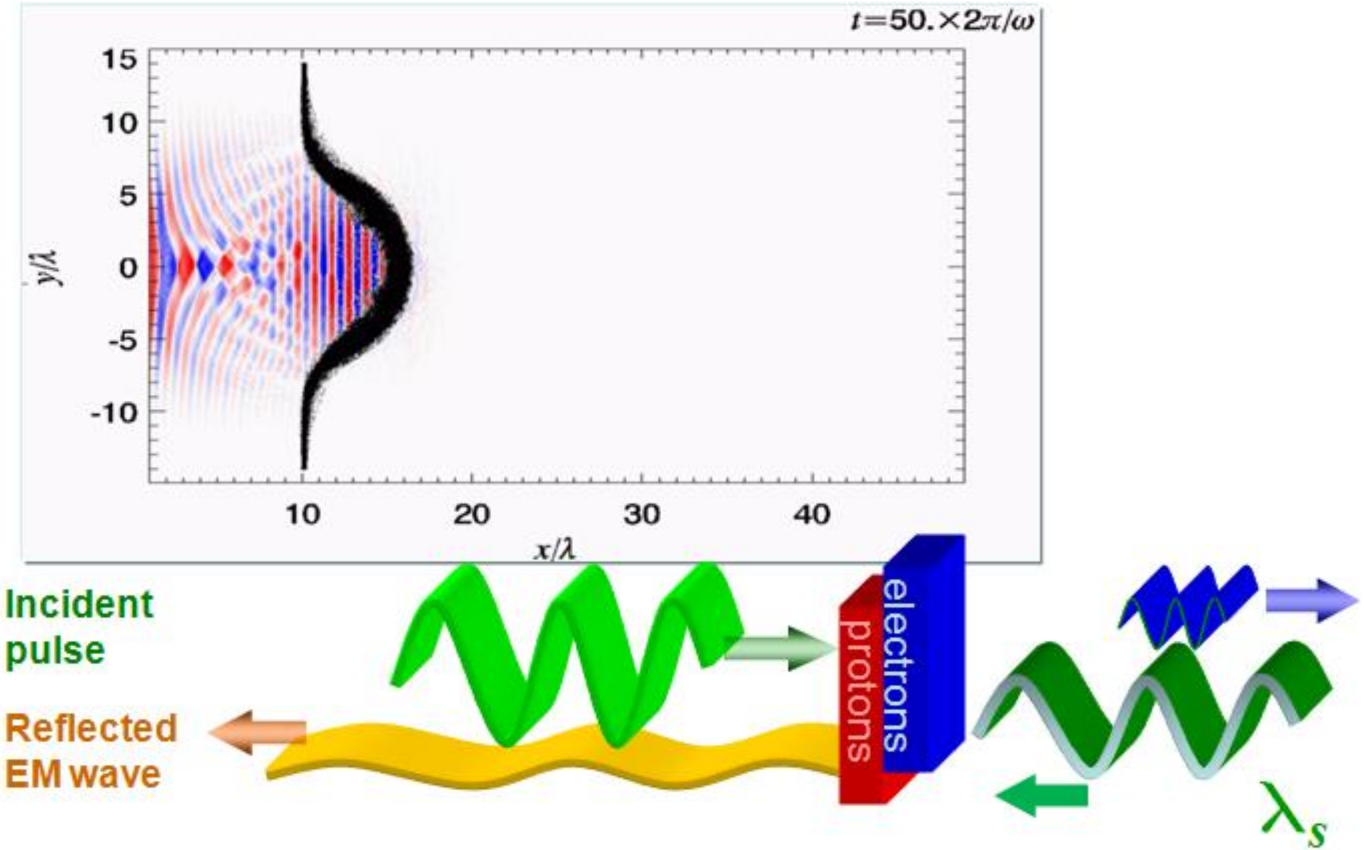




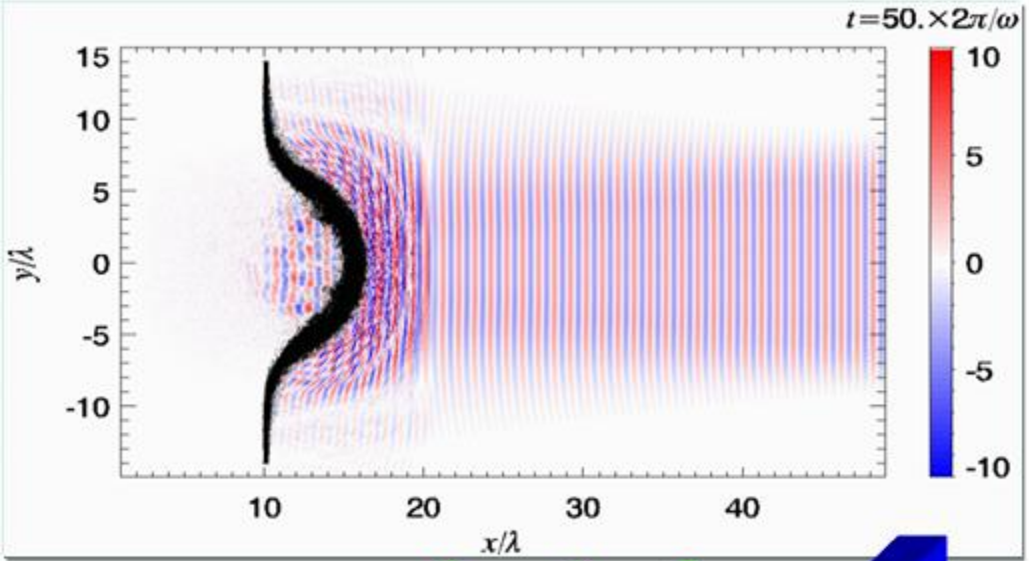


T.Esirkepov, M.Borghesi, SVB, G.Mourou, T.Tajima (2004)
LP or RPDA

Accelerating mirror

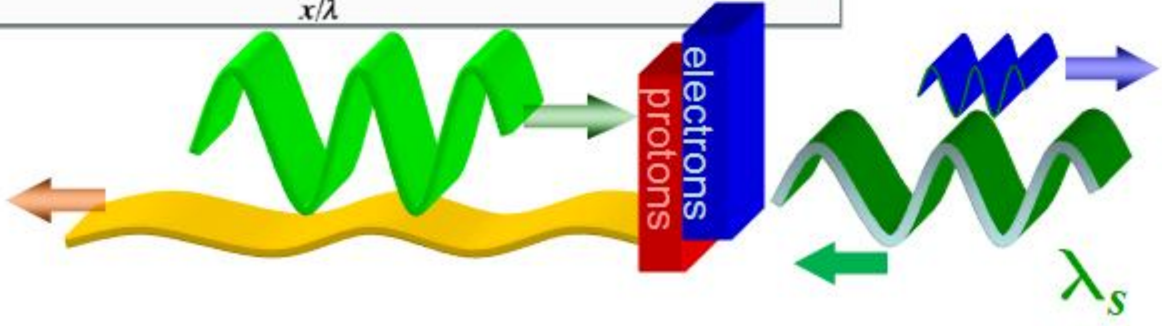


Accelerating mirror



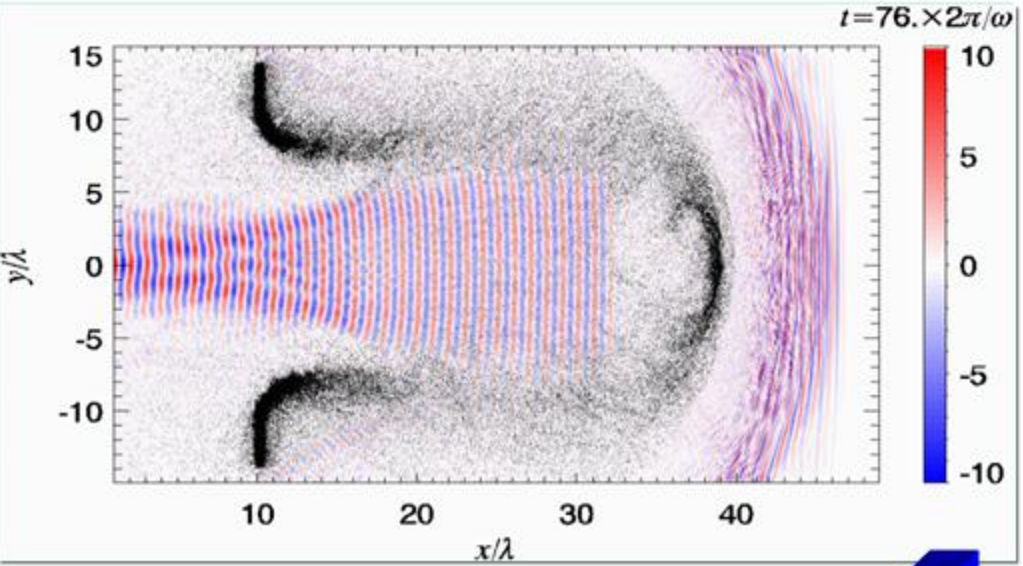
Incident pulse

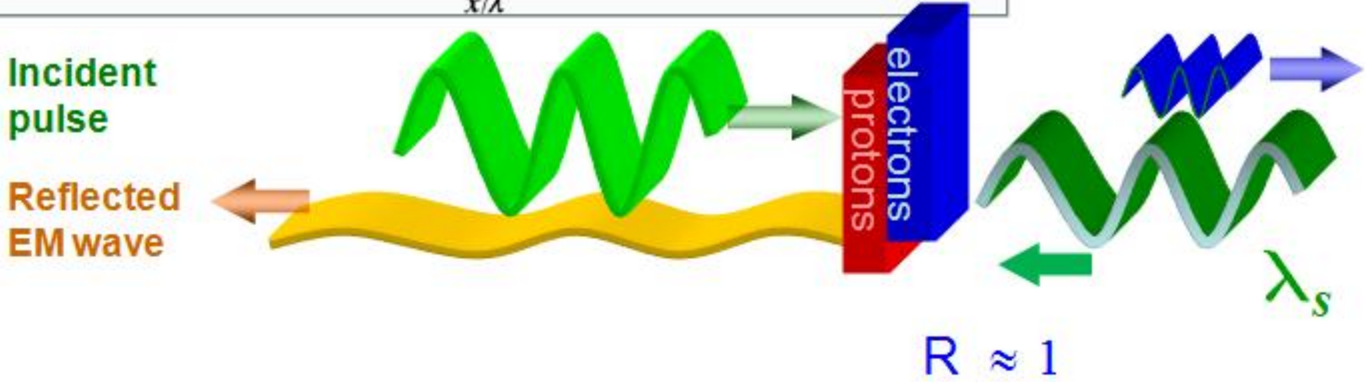
Reflected EM wave



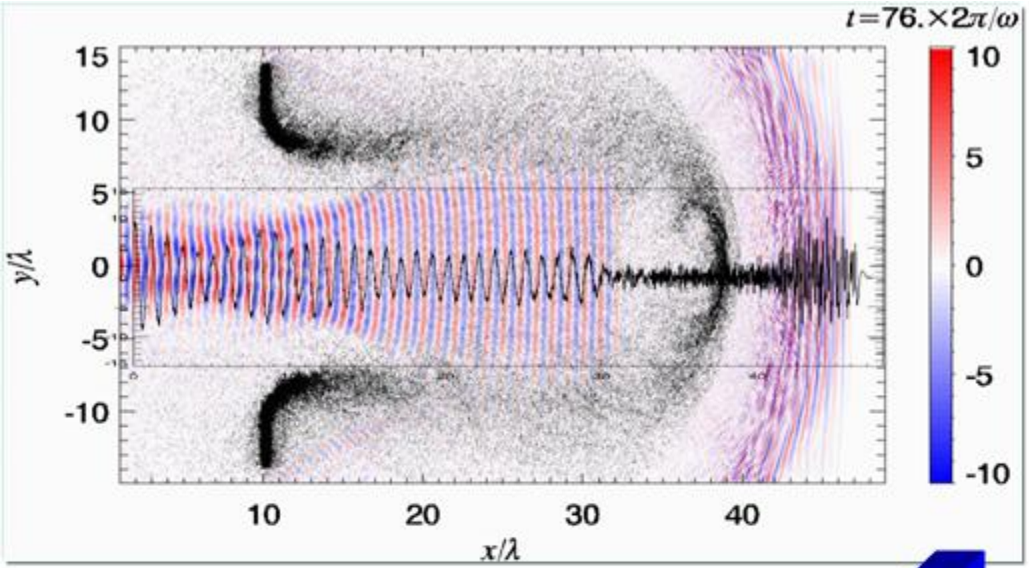
$$\tilde{\lambda}_r \approx \frac{\lambda_s}{4\gamma^2}$$

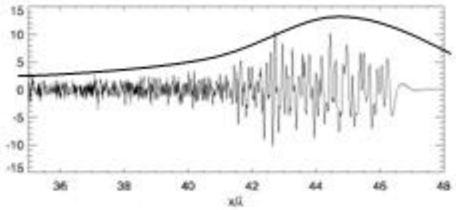
Accelerating mirror

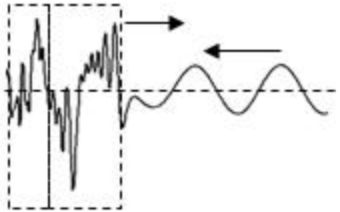


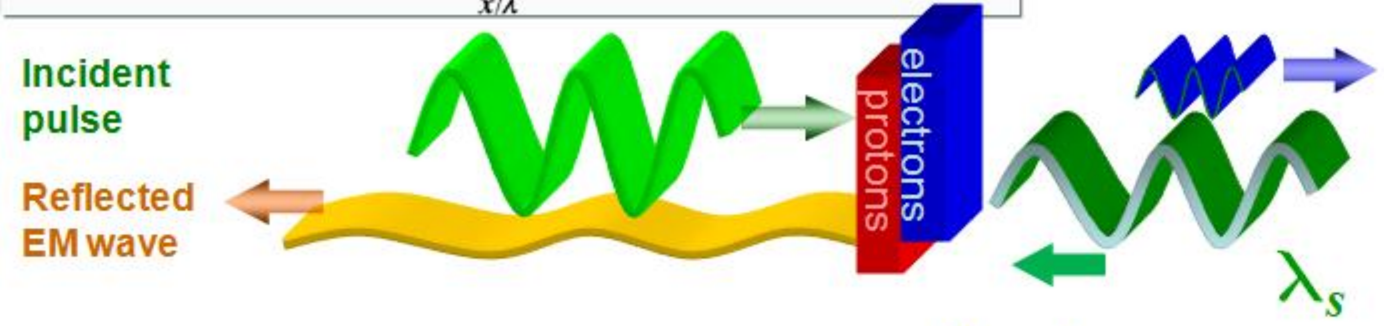


Laser Piston+Flying Mirror+Oscillating Mirror









$$\omega_n = n\omega \cdot 4\gamma^2$$

$$R \approx 1$$

I did not discuss

Weibel Instability
Unipolar Accelerator
Wake Field Acceleration
Buneman Instability…..
Several other realizations of relativistic mirrors

7. Conclusion

a) **Ultra Short Pulse Laser – Matter Interaction has entered the Ultrarelativistic Regime. By this it has opened a new field of Relativistic Laboratory Astrophysics**

b) **Laser Piston+Flying Mirror+Oscillating Mirror will provide in a nearest future the instruments for nonlinear vacuum probing and for studying other fundamental problems**

Xie Xie

Second International Symposium on Laser-Driven Relativistic Plasmas Applied to Science, Industry and Medicine

Kansai Photon Science Institute, JAEA,
Kizugawa, Kyoto, Japan, January 19 to 23, 2009

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a) Charged Particle Acceleration
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c) Relativistic Laser Astrophysics
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